Dust temperature fluctuations and surface chemistry: H_2 formation

Emeric Bron

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Four levels of modeling Dust temperature fluctuations

Four levels of modeling

• Parametrization : $R_{\rm H_2} = R_f n_{\rm H} n({\rm H})$

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Microphysical model

Four levels of modeling Dust temperature fluctuations

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 - \rightarrow rate equations

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 \rightarrow out-of-equilibrium chemistry !

Four levels of modeling Dust temperature fluctuations

Formation mechanisms

Langmuir-Hinshelwood mechanism:



- Van der Waal binding (658 K)
- easy migration (thermal hopping and tunneling, 510 K)
- thermal desorption
- experiments vs. observations

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Formation mechanisms

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- Van der Waal binding (658 K)
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- experiments vs. observations

Eley-Rideal mechanism:



- strong bond (\sim 10000 K)
- no migration
- chemisorption barrier (300 K)
- can explain formation in PDRs (Le Bourlot 2012)

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Dust temperature fluctuations



Dust grains absorb UV photons, re-emit in IR

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- Small grains \rightarrow small heat capacities \rightarrow large temperature fluctuations

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- IR emission : dominant during temperature peaks (Draine & Li 2001)
- How does it affect H₂ formation?



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Master equation approach Approximation method



• Goal: average formation rate

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Master equation approach Approximation method

Method

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• interested in
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Master equation approach Approximation method

Approximation method

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Effects Langmuir-Hinshelwood Eley-Rideal Interstellar cloud simulations

Distribution effects vs. out-of-equilibrium effects



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Langmuir-Hinshelwood



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Langmuir-Hinshelwood



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Effects Langmuir-Hinshelwood Eley-Rideal Interstellar cloud simulations

Langmuir-Hinshelwood



- Strongly enhanced efficiency on small grains (up to 20 nm)
- Dominant effect : spread of the temperature PDF
- Langmuir-Hinshelwood efficient in cloud-edge conditions

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Effects Langmuir-Hinshelwood Eley-Rideal Interstellar cloud simulations

Eley-Rideal



• Only very small grains affected (below 2 nm)

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Effects Langmuir-Hinshelwood Eley-Rideal Interstellar cloud simulations

Eley-Rideal



- Only very small grains affected (below 2 nm)
- Limited effect on the total formation rate

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Effects Langmuir-Hinshelwood Eley-Rideal Interstellar cloud simulations

Interstellar cloud simulations



• Medium and low UV fields $(G_0 \lesssim a \text{ few 100})$: LH becomes dominant

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Interstellar cloud simulations



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- Strong UV fields : ER dominates

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Interstellar cloud simulations



- Medium and low UV fields ($G_0 \lesssim a \text{ few 100}$): LH becomes dominant
- Strong UV fields : ER dominates
- Limited effect on observable lines



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Interstellar cloud simulations



• Out-of-equilibrium effect: competition fluctuation timescale vs. adsorption timescale

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- core = molecular gas \rightarrow rare H atoms
- additionnal fluctuations : secondary UV photons, cosmic rays

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Interstellar cloud simulations



- Out-of-equilibrium effect: competition fluctuation timescale vs. adsorption timescale
- core = molecular gas \rightarrow rare H atoms
- additionnal fluctuations : secondary UV photons, cosmic rays

Possible applications : maximal molecular fraction in the diffuse ISM, residual atomic fraction in dark cores





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Conclusions

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- Dust temperature fluctuations are important for surface chemistry

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- LH formation efficient in most environments except bright PDRs
- More flexibility for the ER parameters
- Dust temperature fluctuations are important for surface chemistry
- Other surface reactions: effects of secondary UV photons ? cosmic rays ?

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